

Gödel type metrics in three dimensions

Metin Gürses

Received: 1 June 2009 / Accepted: 1 November 2009 / Published online: 16 November 2009
© Springer Science+Business Media, LLC 2009

Abstract We show that the Gödel type metrics in three dimensions with arbitrary two dimensional background space satisfy the Einstein-perfect fluid field equations. We also show that there exists only one first order partial differential equation satisfied by the components of fluid's velocity vector field. We then show that the same metrics solve the field equations of the topologically massive gravity where the two dimensional background geometry is a space of constant negative Gaussian curvature. We discuss the possibility that the Gödel type metrics to solve the Ricci and Cotton flow equations. When the vector field u^μ is a Killing vector field, we came to the conclusion that the stationary Gödel type metrics solve the field equations of the most possible gravitational field equations where the interaction lagrangian is an arbitrary function of the electromagnetic field and the curvature tensors.

Keywords Gödel type metrics · Topologically massive gravity · Einstein's equations in three dimensions · Ricci and Cotton flows · Einstein-perfect fluid solutions

1 Introduction

Three dimensional spacetimes have some interesting properties both geometrically and physically. The Weyl tensor vanishes but the spacetime is in general not conformally flat. For conformal flatness in three dimensions the Cotton tensor is essential. Physically, three dimensional Einstein gravity is dynamically trivial. For this purpose

M. Gürses (✉)
Department of Mathematics, Faculty of Sciences,
Bilkent University, 06800 Ankara, Turkey
e-mail: gurses@fen.bilkent.edu.tr

topologically massive gravity (TMG) theories were introduced [1–3]. See the review of Carlip [4] for more details. There are several attempts to find exact solutions of Einstein and topologically massive gravity field equations in three dimensions [5–8]. In these efforts authors usually start with a specific ansatz for the spacetime metrics. It seems that the Gödel type metrics [9–11] would be very convenient and practical in searching for solutions of the field equations in three dimensions.

In a D dimensional spacetime the Gödel type metrics are defined by

$$g_{\mu\nu} = h_{\mu\nu} - u_\mu u_\nu \quad (1)$$

where $h_{\mu\nu}$ is the metric of a $D - 1$ -dimensional locally Euclidean Einstein space with $h_{\mu\nu} u^\mu = 0$ and u_μ is a unit timelike vector field with $u^\mu = -\frac{1}{u_0} \delta_0^\mu$. We studied these metrics when u_0 is a constant in [9] and when u_0 is not constant in [10]. Although, our approach in these works was independent of the dimension of the spacetime we focused our attention to the cases $D > 3$ in great detail. In these works we obtained exact solutions of various supergravity theories in various dimensions. In [9] since u_0 was considered to be a constant our solutions contain no dilaton field. If u_0 is not a constant it plays the role of the dilaton field. In [10] we found exact solutions of the supergravity theories with dilaton. In [11] we studied the closed timelike curves in Gödel type metrics and showed that when the vector field u_μ is also a Killing vector of the spacetime geometry then there always exist closed timelike or null curves in Gödel type spacetimes.

In this work we shall consider the Gödel type metrics in three dimensions with u_0 constant (or g_{00} is a constant). The case when u_0 is not a constant will be discussed later. There are several interesting properties of the spacetime geometry in three dimensions. In two dimensions since the Ricci tensor is proportional to the metric then the metrics of any three dimensional spacetime is of Gödel type. In three dimensions energy momentum tensor of a Maxwell field of the vector field u_μ is equivalent to the energy momentum tensor of a perfect fluid with stiff equation of state. With these properties we show in Sect. 3 that any Gödel type metrics in three dimensions satisfy the Einstein-Perfect fluid field equations. Using the result of Sect. 3 we show in Sect. 4 that Gödel type metrics in three dimensions satisfy the field equations of the TMG provided that the two dimensional space is a space of constant curvature. We find all possible Gödel type solutions of TMG and show that our previous solution [5] of TMG is a special case. The Ricci flow equations play important role in differential geometry, in particular in the proof of the Poincare Lemma [14, 15]. Since the proof of Perelman [15] there is an increasing interest in Ricci and Cotton flow equations. Such flow equations may also have applications in the Hořava gravity [16, 17]. The Ricci and Cotton tensors for the Gödel type metrics take very simple forms which attracts us to consider the corresponding flow equations. We study the Ricci and Cotton flow equations in Sect. 5. In the last Section we construct a closed tensor algebra which enables the Gödel type metrics to solve the field equations of a most general Lagrange function of metric, Ricci, curvature and the antisymmetric Maxwell tensor field and their covariant derivatives at all order.

2 Gödel type metrics in general relativity

For the sake of completeness we give in this section, a short review of Gödel type metrics in general relativity in four and higher dimensions.

Let $u^\mu = -\frac{1}{u_0} \delta_0^\mu$ be a timelike vector with $u_0 = \text{constant}$, in D dimensional spacetime M and $h_{\mu\nu}$ be the metric of $D - 1$ dimensional Euclidean space such that $u^\mu h_{\mu\nu} = 0$. Gödel type of metrics are defined by [9]

$$g_{\mu\nu} = h_{\mu\nu} - u_\mu u_\nu. \tag{2}$$

Let us define an antisymmetric tensor $f_{\mu\nu}$ as

$$f_{\alpha\beta} = u_{\beta,\alpha} - u_{\alpha,\beta}. \tag{3}$$

The Christoffel symbol corresponding to the metric (1) is

$$\Gamma_{\alpha\beta}^\mu = \gamma_{\alpha\beta}^\mu + \frac{1}{2} (u_\alpha f^\mu{}_\beta + u_\beta f^\mu{}_\alpha) - \frac{1}{2} (u_{\alpha|\beta} + u_{\beta|\alpha}) u^\mu, \tag{4}$$

where a vertical stroke denotes covariant derivative with respect to the Christoffel symbol $\gamma_{\alpha\beta}^\mu$ and a semicolon or nabla ∇ will denote covariant derivative with respect to the Christoffel symbol $\Gamma_{\alpha\beta}^\mu$. It is easy to show that

$$u^\alpha \partial_\alpha u_\beta = 0, \quad u^\alpha f_{\alpha\beta} = 0. \tag{5}$$

Then

$$\dot{u}^\mu = u^\alpha u_{;\alpha}^\mu = 0. \tag{6}$$

It is a property of the Gödel type of geometries that the vector field u^μ is also a timelike Killing vector field of the spacetime geometry (M, g) and hence we have

$$f_{\alpha\beta} = 2 u_{\beta;\alpha}. \tag{7}$$

We define the current vector j_μ corresponding to the antisymmetric tensor field $f_{\mu\nu}$ as

$$j_\mu \equiv f^\alpha{}_{\mu|\alpha} = \nabla_\alpha f^\alpha{}_\nu - \frac{1}{2} f^2 u_\nu, \tag{8}$$

where $f^2 = f_{\mu\nu} f^{\mu\nu}$.

It is now straightforward to prove the following Proposition [9]

Proposition 1 *Let (M, g) be a stationary spacetime with the Gödel type metric (1). Let $h_{\mu\nu}$ be the metric tensor of $D - 1$ -dimensional locally Euclidean space, then the*

Einstein tensor becomes

$$G_{\mu\nu} = r_{\mu\nu} - \frac{1}{2} h_{\mu\nu} r + \frac{1}{2} T_{\mu\nu}^f + \frac{1}{2} (j_\mu u_\nu + j_\nu u_\mu) + \left(\frac{1}{4} f^2 + \frac{1}{2} r \right) u_\mu u_\nu - \frac{1}{2} (u^\alpha j_\alpha) g_{\mu\nu}, \quad (9)$$

where $T_{\mu\nu}^f$ denotes the Maxwell energy-momentum tensor for $f_{\mu\nu}$, $r_{\mu\nu}$ is the Ricci tensor of $\gamma^\mu{}_\alpha\beta$. The Ricci scalar is obtained as

$$R = r + \frac{1}{4} f^2 + u^\mu j_\mu,$$

where r denotes the Ricci scalar of $r_{\alpha\beta}$.

The above Proposition gives the Einstein tensor of (1) without any conditions. In order to have a physical energy momentum distribution we assume that $D - 1$ -dimensional space is an Einstein space (a vacuum space with a cosmological constant) and the current vector field j_μ vanishes everywhere. Then we have

Proposition 2 *Let (M, g) be a stationary spacetime geometry with the Gödel type metric (1). Let $h_{\mu\nu}$ be the metric tensor of the $D - 1$ -dimensional Einstein space, $r_{\mu\nu} = \frac{r}{D-1} h_{\mu\nu}$ and let $j_\mu = 0$. Then the metric $g_{\mu\nu}$ satisfies the Einstein field equations with a charged fluid*

$$G_{\mu\nu} = \frac{1}{2} T_{\mu\nu}^f + (p + \rho) u_\mu u_\nu + p g_{\mu\nu}, \quad (10)$$

with

$$\nabla_\mu f^{\mu\nu} = \frac{1}{2} f^2 u^\nu, \quad (11)$$

$$p = \frac{(3 - D)}{2(D - 1)} r, \quad (12)$$

$$\rho = \frac{1}{4} f^2 + r. \quad (13)$$

Here p is the pressure and ρ is the energy density of the charged perfect fluid.

Here the signs of the fluid pressure and the fluid energy density depends (in particular for $D \geq 3$) on the sign of the Ricci scalar of the $D - 1$ dimensional Euclidean space with metric $h_{\mu\nu}$.

Corollary 3 *If $h_{\mu\nu}$ is the metric of a Ricci flat space then the energy momentum distribution for the Einstein field equations for the metric $g_{\mu\nu}$ becomes charged dust, i.e.,*

$$G_{\mu\nu} = \frac{1}{2} T_{\mu\nu}^f + \frac{1}{4} f^2 u_\mu u_\nu \quad (14)$$

provided f satisfies the equation

$$f^\alpha{}_{\beta|\alpha} = 0 \tag{15}$$

and T^f is the Maxwell energy momentum tensor for the antisymmetric tensor f

$$T^f_{\mu\nu} = f_{\mu\alpha} f_\nu{}^\alpha - \frac{1}{4} f^2 g_{\mu\nu}$$

where $f^2 = f^{\alpha\beta} f_{\alpha\beta}$. Maxwell's equation (15) can also be written as (11).

Hence Gödel type metrics (1) satisfy the Einstein field equations with charged dust distributions where the only field equations are the Maxwell equations (15) or (11) and the Ricci flat equations for $h_{\mu\nu}$. There is no electric field ($u^\mu f_{\mu i} = f_{0i} = 0$), only the magnetic field exists. We have the gauge freedom

$$u'{}_\mu = u_\mu + \partial_\mu \Omega \tag{16}$$

where Ω is a function satisfying the condition (recall that $u'^{\mu} = -\frac{\delta_0^\mu}{u'_0}$, $u^\mu = -\frac{\delta_0^\mu}{u_0}$)

$$u'_0 = u_0 + \partial_0 \Omega \tag{17}$$

Both u_0 and u'_0 are constants. For the stationary spacetime, which is the case in this work, we have Ω not depending on x^0 and $u'_0 = u_0$ (constant) but this leads to constant Ω .

$$u'^{\mu} = u^\mu + g^{\mu\nu} \Omega_{, \nu} \tag{18}$$

or

$$\frac{\delta_0^\mu}{u'_0} = \frac{\delta_0^\mu}{u_0} - g^{\mu\nu} \Omega_{, \nu} \tag{19}$$

since $u'_0 = u_0$ then $\Omega = constant$. Hence we have the following Proposition.

Proposition 4 *The only gauge transformation*

$$h'_{\mu\nu} = h_{\mu\nu}, \quad u'{}_\mu = u_\mu + \partial_\mu \Omega \tag{20}$$

keeping the stationary Gödel type metric invariant is the one with constant Ω .

In this work we only considered the case where u_0 is a non vanishing constant. In [10] we have studied the Gödel type metrics when u_0 is not a constant. In this case, the proposed metric yields exact solutions to the various theories with a dilaton field.

3 Gödel type metrics in three dimensions

In three dimensions Gödel type of metrics have very interesting properties. All three dimensional metrics can be written as a Gödel type of metric with a non-constant u_0 .

Proposition 5 *In three dimensions all metrics are of Gödel type.*

Proof Any three dimensional metric can be written as follows

$$\begin{aligned} ds^2 &= -P^2 (dx^0)^2 + 2Mdx^0 dx^1 + 2Ndx^0 dx^2 \\ &\quad + Q^2 (dx^1)^2 + 2Ldx^1 dx^2 + R^2 (dx^2)^2 \\ &= -P^2 \left(dx^0 - \frac{M}{P^2} dx^1 - \frac{N}{P^2} dx^2 \right)^2 + \left(\frac{M^2}{P^2} + Q^2 \right) (dx^1)^2 \\ &\quad + 2 \left(L + \frac{MN}{P^2} \right) dx^1 dx^2 + \left(\frac{N^2}{P^2} + R^2 \right) (dx^2)^2, \end{aligned} \quad (21)$$

where P, M, N, Q, L, R are functions of x^0, x^1 and x^2 . Then the last form (21) is of Gödel type (1) with

$$\begin{aligned} u_\mu dx^\mu &= P \left(dx^0 - \frac{M}{P^2} dx^1 - \frac{N}{P^2} dx^2 \right), \\ h_{\mu\nu} dx^\mu dx^\nu &= \left(\frac{M^2}{P^2} + Q^2 \right) (dx^1)^2 + 2 \left(L + \frac{MN}{P^2} \right) dx^1 dx^2 \\ &\quad + \left(\frac{N^2}{P^2} + R^2 \right) (dx^2)^2 \end{aligned} \quad (22)$$

Hence $u_0 = P$ which is not a constant in general and h is the metric of a two dimensional locally Euclidean space.

Corollary 6 *When (M, g) is stationary and $u_0 = \text{constant}$ then metric functions depend on x^1 and x^2 , and $P = \text{constant}$*

Another interesting property of three dimensions is that any antisymmetric tensor field $f_{\mu\nu}$ can be expressed as $\epsilon_{\mu\nu\alpha} v^\alpha$ where v^α is any vector field. Since u_0 is constant and $u_{\alpha,0} = 0$ (stationarity) then $u^\alpha f_{\alpha\beta} = 0$. This implies that v^α is proportional to u^α . Hence we have the following Proposition:

Proposition 7 *The antisymmetric tensor $f_{\mu\nu}$ can be expressed as*

$$f_{\mu\nu} = 2w \eta_{\mu\nu\alpha} u^\alpha \quad (23)$$

where w is an arbitrary function and $\eta_{\mu\nu\alpha} = \sqrt{|g|} \epsilon_{\mu\nu\alpha}$. Here $\epsilon_{\mu\nu\alpha}$ is the totally antisymmetric Levi-Civita tensor. Hence from (7) we have

$$\nabla_\mu u_\nu = w \eta_{\mu\nu\alpha} u^\alpha \quad (24)$$

Taking the divergence of (23) we obtain

$$\nabla_\mu f^{\mu\nu} = \frac{1}{2} f^2 u^\nu + 2w_{,\alpha} \eta^{\alpha\nu\mu} u_\mu, \tag{25}$$

where $w^2 = \frac{1}{8} f^2$. This leads to the following result:

Proposition 8 *The above Eq. (25) implies that in three dimensions the Maxwell equations (11) are satisfied if and only if w or $f^2 = \text{constant}$.*

In three dimensions due to the property (23) the energy momentum tensor of $f_{\mu\nu}$ becomes the energy momentum tensor of a perfect fluid with $p = \rho$ equation of state.

Proposition 9 *In three dimensions due to the property (23) the energy momentum tensor corresponding to the antisymmetric tensor field $f_{\mu\nu}$ reduces to*

$$T^f_{\mu\nu} = \frac{1}{2} f^2 u_\mu u_\nu + \frac{1}{4} f^2 g_{\mu\nu}, \tag{26}$$

where the energy density and the pressure are respectively given by

$$\rho = \frac{1}{4} f^2, \quad p = \frac{1}{4} f^2. \tag{27}$$

Hence we have the stiff equation of state $p = \rho$.

Then any stationary spacetime metric in three dimensions with $f^2 = \text{constant}$ satisfies the Einstein perfect fluid field equations. We state this as the next proposition which will be used later for different purposes.

Proposition 10 *Let $D = 3$ in Proposition 2 and use Proposition 9 for the energy momentum tensor of $f_{\mu\nu}$ then the stationary Gödel type metrics (1) with constant f^2 satisfy the Einstein field equation with a perfect fluid distribution*

$$G_{\mu\nu} = \frac{1}{2} (f^2 + r_2) u_\mu u_\nu + \frac{1}{8} f^2 g_{\mu\nu}, \tag{28}$$

where r_2 is the Ricci scalar corresponding to the two dimensional metric tensor $h_{\mu\nu}$. Energy density and the pressure of the fluid are respectively given by

$$p = \frac{1}{8} f^2, \quad \rho = \frac{3}{8} f^2 + \frac{1}{2} r_2 \tag{29}$$

To have some specific solutions we need a coordinate chart. For this purpose let us now consider the metric in polar coordinates (geodesic polar coordinates).

Proposition 11 *Without losing any generality we can write the metric given in (21) in polar coordinates so that the two dimensional part the coordinate curves are orthogonal*

$$ds^2 = m^2 dr^2 + n^2 d\theta^2 - (u_0 dt + u_1 dr + u_2 d\theta)^2, \tag{30}$$

where $u_\mu = (u_0, u_1, u_2)$. Here u_0 is a constant, $x^\mu = (t, r, \theta)$, the functions m, n, u_1 and u_2 depend on r and θ . The only field equation (23) reduces to a single equation

$$u_{1,\theta} = u_{2,r} + 2wmn, \quad (31)$$

which is equivalent to $f^2 = 8w^2 = \text{constant}$. As a conclusion, for any 3 dimensional metric $g_{\mu\nu}$ with g_{00} is a constant, Eq. (31) solves the stationary Einstein-Perfect fluid equations where the pressure and energy density are given in (29).

4 Gödel type metrics in topologically massive theory

Topologically massive gravity (TMG) equations found by Deser, Jackiw and Templeton (DJT) [1, 2] with a cosmological constant are given as follows.

$$G_\nu^\mu + \frac{1}{\mu} C_\nu^\mu = \lambda \delta_\nu^\mu. \quad (32)$$

Here $G_{\mu\nu}$ and $R_{\mu\nu}$ are the Einstein and Ricci tensors respectively and C_ν^μ is the Cotton tensor which is given by

$$C_\nu^\mu = \eta^{\mu\beta\alpha} \nabla_\alpha \left(R_{\nu\beta} - \frac{1}{4} R g_{\nu\beta} \right). \quad (33)$$

The constants μ and λ are respectively the DJT parameter and the cosmological constant.

We had introduced a method [5] to solve the DJT field equations. In this method we start with Einstein's equations with a perfect fluid source

$$G_{\mu\nu} = T_{\mu\nu}, \quad (34)$$

with

$$T_{\mu\nu} = (p + \rho) u_\mu u_\nu + p g_{\mu\nu}, \quad (35)$$

where the fluid equations are obtained through the conservation equation $\nabla_\mu T^{\mu\nu} = 0$, p and ρ are respectively the pressure and energy density of the fluid and u_μ is the fluid's timelike unit four velocity vector, i.e., $u^\mu u_\mu = -1$. We have the following result which was reported previously [5].

Proposition 12 *If p, ρ are constants and*

$$\nabla_\mu u_\nu = \frac{\mu}{3} \eta_{\mu\nu\alpha} u^\alpha \quad (36)$$

then any solution of the Einstein equations $G_{\mu\nu} = T_{\mu\nu}$ with a perfect fluid distribution is also a solution of the TMG (32) with a cosmological constant $\lambda = \frac{2p-\rho}{3}$ and

$$p = \frac{\mu^2}{9}, \quad \rho = \frac{\mu^2 - 9\lambda_0}{3} \tag{37}$$

where $\lambda_0 = \lambda + \frac{\mu^2}{27}$.

If the Ricci scalar r_2 of the metric $h_{\mu\nu}$ is a constant then as a consequence of the Proposition 12 any Gödel type of metrics solve the DJT equations

Proposition 13 Stationary Gödel type metrics in three dimensions with constant f^2 , see Proposition 10, solve also the TMG field equations if the two dimensional background space is of constant Gaussian curvature (or $r_2 = \text{constant}$) and

$$\mu = 3w, \quad r_2 = -2(w^2 + 3\lambda) \tag{38}$$

Equation (38) implies that the two dimensional geometry with the metric $h_{\mu\nu}$ is flat if $\lambda_0 = 0$. As an application of the above Proposition 13 let us consider the following solution of the TMG [5].

Proposition 14 The following metric solves the TMG equations exactly

$$\begin{aligned} ds^2 &= -a_0 dt^2 + 2q dt d\theta + \frac{-q^2 + h^2 \psi}{a_0} d\theta^2 + \frac{1}{\psi} dr^2 \\ &= -\left(\sqrt{a_0}dt - \frac{q}{\sqrt{a_0}} d\theta\right)^2 + \frac{h^2 \psi}{a_0} d\theta^2 + \frac{1}{\psi} dr^2 \end{aligned} \tag{39}$$

where $u_0 = \sqrt{a_0}$, $u_1 = 0$ and $u_2 = -\frac{q}{\sqrt{a_0}}$ and

$$\begin{aligned} \psi &= b_0 + \frac{b_1}{r^2} + \frac{3\lambda_0}{4} r^2 \\ q &= c_0 + \frac{e_0 \mu}{3} r^2 \\ h &= e_0 r, \quad \lambda_0 = \lambda + \frac{\mu^2}{27} \end{aligned}$$

where a_0, b_0, b_1, c_0 and e_0 are arbitrary constants.

The above solution is stationary and has radial symmetry. All the metric functions depend only on the radial coordinate r . We may call this solution as stationary spherically symmetric (solutions not depending on the angular coordinate θ) Gödel type metrics.

Remark 1 The properties of the above metric (39) is discussed in [5]. Although there are five constants in this metric some of them are redundant. They are reduced to the physical constants mass M and angular momentum J in [5].

Remark 2 In the study of the black holes solutions in TMG, Moussa et al. [6] considered the solution of Vuorio [12, 13] which is given by

$$ds^2 = -[d\tilde{t} - (2 \cosh \sigma + \tilde{w})d\tilde{\varphi}]^2 + d\sigma^2 + \sinh^2 \sigma d\tilde{\varphi}^2. \quad (40)$$

where \tilde{w} is a constant. This solution is a special case of our solution (39) with $\lambda = 0$, $\mu = 3$ and

$$b_0 = -2, \quad b_1 = 0, \quad e_0 = 1/2, \quad a_0 = 1, \quad c_0 = \tilde{w} - 2 \quad (41)$$

The transformation links our solution to the Vuorio solution is given as $t = \tilde{t}$, $\theta = \tilde{\varphi}$, $r^2 = 4(1 + \cosh \sigma)$. Our analysis shows that the Vuorio solution is also of Gödel type. Hence one may use a similar analytic continuation used in [6] to our solution (39) with $\lambda \neq 0$, $\mu \neq 3$ to convert it to a black hole solution of TMG. We remark (Clement, private communication) also that the solution given in (39) is equivalent to the solution (3.13) of Moussa et al. [7].

We shall now give, using Proposition 10, a generalization of the above spherically symmetric solution of TMG. We shall do this by taking the metric of the two dimensional geometry same as the one given in (39)

$$ds_2^2 = \frac{h^2 \psi}{a_0} d\theta^2 + \frac{1}{\psi} dr^2 \quad (42)$$

where ψ and h are given in Proposition 14. This leads to a constant Gaussian curvature $K = 3\lambda_0 = 3\lambda + w^2$. Hence

$$m = \frac{1}{\sqrt{\psi}}, \quad n = \frac{h\sqrt{\psi}}{\sqrt{a_0}} \quad (43)$$

and take the most general solution of Eq. (31), i.e.,

$$u_{1,\theta} = u_{2,r} + 2w \frac{e_0}{\sqrt{a_0}} r \quad (44)$$

We solve u_2 from this equation as

$$u_2 = -\frac{1}{\sqrt{a_0}}(c_0 + we_0 r^2) + \int^r (u_{1,\theta}) dr \quad (45)$$

where c_0 is arbitrary constant and u_1 is left free. Hence the metric

$$ds^2 = m^2 dr^2 + n^2 d\theta^2 - (u_0 dt + u_1 dr + u_2 d\theta)^2 \quad (46)$$

with m , n , and u_2 given above (43) and (45) respectively, solve the TMG exactly. Here $u_1(r, \theta)$ is left arbitrary which was taken to be zero in our solution (39). The following Proposition summarizes this result.

Proposition 15 *We obtain the most general stationary solution of TMG when g_{00} is a constant. The solution is given in Gödel type where $u_0 = \sqrt{a_0}$, u_1 is an arbitrary function of r and θ , u_2 is given in (45) and the two dimensional metric is given in (42) with constant Gaussian curvature $K = 3\lambda_0$. This solution generalizes our solution presented in [5].*

5 Ricci and Cotton flows

In this section we shall assume that the Gaussian curvature K of the two dimensional space with metric $h_{\mu\nu}$ is a constant. From Proposition 10 we have the Ricci tensor of a stationary Gödel type metrics

$$R_{\mu\nu} = \frac{1}{2} (f^2 + r_2) u_\mu u_\nu + \frac{1}{2} \left(r_2 + \frac{1}{2} f^2 \right) g_{\mu\nu} \tag{47}$$

where $r_2 = 2K$ is the Ricci scalar corresponding to the two dimensional metric tensor $h_{\mu\nu}$. Then we have an exact solution of the Ricci flow equation

Proposition 16 *Let (M, g) be the stationary Gödel type spacetime with $f^2 = 8w^2$ constant as in Proposition 10. Then Ricci flow equation [14]*

$$\frac{\partial g_{\mu\nu}}{\partial s} = \xi R_{\mu\nu} \tag{48}$$

where s is the flow parameter and ξ is an arbitrary constant, has an exact solution if

$$\frac{\partial u_\mu}{\partial s} = -\xi w^2 u_\mu, \tag{49}$$

$$\frac{\partial h_{\mu\nu}}{\partial s} = -\xi (K - 2w^2) h_{\mu\nu}. \tag{50}$$

The above flow equations (49) and (50) are solved exactly by playing with the constants b_0, c_0, λ_0 , and e_0 . As an example u_0 (or a_0) has the following behavior under the this flow

$$u_0 = u_0(0) e^{-\xi w^2 s} \tag{51}$$

where $u_0(0)$ is an arbitrary constant. On the other hand taking the trace of both sides of (48) we obtain that $u_0 = u_0(0) e^{-\xi(w^2+K)s}$. Hence comparing with (51) we get $K = 0$.

Proposition 17 *The Ricci flow equations have Gödel type of metrics as exact solutions only when the two dimensional space is a space of zero curvature*

Proposition 18 *The Cotton tensor for stationary Gödel type metrics take the following simple form*

$$C^\mu{}_\nu = -w (p + \rho) (\delta^\mu_\nu + 3u^\mu u_\nu). \tag{52}$$

Hence the spacetime geometry (M, g) is conformally flat if $\lambda = \frac{\mu^2}{9}$. The Cotton flow equation

$$\frac{\partial g_{\mu\nu}}{\partial s} = \zeta C_{\mu\nu} \quad (53)$$

where ζ is a constant and s is the flow parameter. These equations were recently used by [18]. Here we show that Gödel type metrics solve exactly the Cotton flow equations (53) only when the Cotton tensor vanishes. For $C_{\mu\nu}$ given above we have

Proposition 19 *Let (M, g) be the stationary Gödel type spacetime with $f^2 = 8w^2$ constant as in Proposition 10. Then Cotton flow equations (53) lead to the following flow equations for $h_{\mu\nu}$ and u_μ*

$$\frac{\partial u_\mu}{\partial s} = -k u_\mu, \quad (54)$$

$$\frac{\partial h_{\mu\nu}}{\partial s} = -k h_{\mu\nu}. \quad (55)$$

where $k = \zeta w (p + \rho)$.

Again the above flow equations (54) and (55) are solved exactly by playing with the constants b_0 , c_0 , λ_0 , and e_0 . As an example u_0 (or a_0) has the following behavior under the Cotton flow

$$u_0 = u_0(0) e^{-ks} \quad (56)$$

where $u_0(0)$ is an arbitrary constant. On the other hand taking the trace of both sides of (53) we obtain u_0 is also constant with respect to the flow parameter s . Comparing this with (56) we obtain $k = 0$.

Proposition 20 *Let (M, g) be the stationary Gödel type spacetime with $f^2 = 8w^2$ constant as in Proposition 10. Then Cotton flow equations (53) have exact solutions only when the $k = 0$, but this means that the Cotton tensor vanishes.*

6 Gödel type metrics in higher curvature theories

In three dimensions when the stationary Gödel type metrics with constant f^2 have further nice properties. It is possible to show that the tensors u_μ , $f_{\mu\nu}$ and $g_{\mu\nu}$ satisfy the following tensorial algebra.

Proposition 21 *Let $D = 3$ and the metric of spacetime be Gödel type with constant f^2 . Let the two dimensional space with metric $h_{\mu\nu}$ be a space of constant Gaussian curvature. Then we have the following closed differential algebra of the tensors*

$u_\mu, f_{\mu\nu}$ and $g_{\mu\nu}$

$$\nabla_\mu u_\alpha = \frac{1}{2} f_{\mu\alpha}, \tag{57}$$

$$f_{\mu\alpha} = w \epsilon_{\mu\alpha\sigma} u^\sigma, \tag{58}$$

$$\nabla_\alpha f_{\mu\beta} = \frac{1}{2} w^2 (g_{\mu\alpha} u_\beta - g_{\beta\alpha} u_\mu), \tag{59}$$

$$\nabla_\mu g_{\alpha\beta} = 0. \tag{60}$$

From the previous Propositions (in particular Proposition 10) we can deduce that The Ricci, Einstein curvature tensors and their contractions at any order will be the linear sum the tensors $u_\mu u_\nu$ and $g_{\mu\nu}$. Hence the tensor differential algebra introduced in (57)–(60) is effective to show that the gravitational field equations, for any gravitational action. are given as follows

$$G_{\mu\nu} = Ag_{\mu\nu} + Bu_\mu u_\nu, \tag{61}$$

$$\nabla_\mu f^{\mu\nu} = Cu^\nu \tag{62}$$

where A, B and C are constants depending upon the theory. This leads to the following result:

Proposition 22 *Let the action of gravitation contains all possible combinations of Ricci, curvature and the antisymmetric tensor $F_{\mu\nu}$ and their covariant derivatives at any order. Then the tensor differential algebra introduced in (57)–(60) is effective to show that the gravitational field equations are solved when the metric is the stationary Gödel type metrics with constant f^2 and the two dimensional background is a space of constant Gaussian curvature K , and $F_{\mu\nu} = f_{\mu\nu}$ at all orders of the string tension parameter.*

7 Conclusion

We showed that the metric of any three dimensional stationary spacetime with g_{00} constant satisfies the Einstein-perfect fluid equations. The only differential equation to be solved is a first order partial differential for the components of the fluid velocity vector field. We then showed that in this spacetime symmetry with g_{00} constant we find the most general solution of the TMG. This solution generalizes our previous solution [5] (Proposition 15). We showed that stationary Gödel type metrics constitute a very simple solution of the Ricci flow equations an do not solve the Cotton flow equations (Propositions 19 and 20). Finally we discussed the possibility that the stationary Gödel type metrics form a solution of the low energy limit of string theory with the most possible interactions of curvature and antisymmetric field $F_{\mu\nu}$ (Proposition 22).

Acknowledgments I would like to thank Professors Gerard Clement and Atalay Karasu for reading the manuscript and constructive comments. This work is partially supported by the Scientific and Technological Research Council of Turkey (TUBITAK) and Turkish Academy of Sciences (TUBA).

References

1. Deser, S., Jackiw, R., Templeton, S.: *Phys. Rev. Lett.* **48**, 975 (1982)
2. Deser, S., Jackiw, R., Templeton, S.: *Ann. Phys. NY* **140**, 372 (1982)
3. Bergshoeff, E.A., Holm, D., Townsend, P.K.: *Phys. Rev. Lett.* **102**, 201301 (2009)
4. Carlip, S.: *Class. Quantum Grav.* **12**, 2853–2880 (1995)
5. Gürses, M.: *Class. Quantum Grav.* **11**, 2585 (1994)
6. Moussa, K.A., Clement, G., Leygnac, C.: *Class. Quantum Grav.* **20**, L277–L283 (2003)
7. Moussa, K.A., Clement, G., Guennoune, H., Leygnac, C.: *Phys. Rev. D* **78**, 064065 (2008)
8. Barrow, J.D., Shaw, D.J., Tsagas, C.G.: *Class. Quantum Grav.* **23**, 5291–5322 (2006)
9. Gürses, M., Karasu, A., Sarioğlu, Ö.: *Class. Quantum Grav.* **22**, 1527–1543 (2005)
10. Gürses, M., Sarioğlu, Ö.: *Class. Quantum Grav.* **22**, 4699 (2005)
11. Gleiser, R.J., Gürses, M., Karasu, A., Sarioğlu, Ö.: *Class. Quantum Grav.* **23**, 2653 (2006)
12. Vurorio, I.: *Phys. Lett. B* **163**, 91 (1985)
13. Percacci, R., Sodano, P., Vurorio, I.: *Ann. Phys. NY* **176**, 344 (1987)
14. Hamilton, R.S.: *J. Differ. Geom.* **17**, 255 (1982)
15. Perelman, G.: Finite extinction time for the solutions to the Ricci flow on certain three-manifolds. [arXiv:math/0307245v1](https://arxiv.org/abs/math/0307245v1) [math.DG]
16. Hořova, P.: *Phys. Rev. Lett.* **102**, 23130 (2009)
17. Lü, H., Mei, L., Pope, C.N.: *Phys. Rev. Lett.* **103**, 091301 (2009)
18. Kışisel, A.U., Sarioğlu, Ö., Tekin, B.: *Class. Quantum Grav.* **25**, 165019 (2008)